

# Using SDO's AIA to investigate energy transport from a flare's energy release site to the chromosphere

J. W. Brosius<sup>1</sup> and G. D. Holman<sup>2</sup>

<sup>1</sup> The Catholic University of America at NASA's Goddard Space Flight Center, Code 671, Greenbelt, MD 20771 USA  
e-mail: Jeffrey.W.Brosius@nasa.gov

<sup>2</sup> Code 671, NASA's Goddard Space Flight Center, Greenbelt, MD 20771 USA  
e-mail: Gordon.D.Holman@nasa.gov

Received 23 September 2011; Accepted 25 November 2011

## ABSTRACT

**Context.** Coordinated observations of a *GOES* B4.8 microflare with *SDO*'s Atmospheric Imaging Assembly (AIA) and the Ramaty High Energy Solar Spectroscopic Imager (*RHESSI*) on 2010 July 31 show that emission in all seven of AIA's EUV channels brightened simultaneously nearly 6 minutes before *RHESSI* or *GOES* detected emission from plasma at temperatures around 10 MK.

**Aims.** To help interpret these and AIA flare observations in general, we characterized the expected temporal responses of AIA's 94, 131, 171, 193, 211, and 335 Å channels to solar flare brightenings by combining (1) AIA's nominal temperature response functions available through SSWIDL (Boerner *et al.*) with (2) EUV spectral line data observed in a flare loop footpoint on 2001 April 24 with the Coronal Diagnostic Spectrometer (CDS) on timescales comparable to AIA's image cadence (Brosius & Phillips).

**Methods.** The nine emission lines observed by CDS cover a wide range of formation temperature from about 0.05 to 8 MK. Line brightenings observed early during the CDS flare occurred at temperatures less than about 0.7 MK, with the largest values around 0.1 MK. These brightenings were consistent with the flare's energy transport being dominated by nonthermal particle beams. Because all of AIA's EUV channels are sensitive to emission from plasma in the 0.1 to 0.7 MK temperature range, we show that all of AIA's EUV channels will brighten simultaneously during flares like this, in which energy transport is dominated by nonthermal particle beams.

**Results.** The 2010 July 31 flare observed by AIA and *RHESSI* displays this behavior, so we conclude that such beams likely dominated the flare's energy transport early during the event. When thermal conduction from a reconnection-heated, hot (~ 10 MK) plasma dominates the energy transport, the AIA channels that are sensitive to emission from such temperatures (particularly the 94 and 131 Å channels) will brighten earlier than the channels that are not sensitive to such temperatures (171 and 211 Å).

**Conclusions.** Thus, based on the differences expected between AIA's response to flares whose energy transport is dominated by nonthermal particle beams from those whose energy transport is dominated by thermal conduction, AIA can be used to determine the dominant energy transport mechanism for any given event.

**Key words.** Sun: activity – Sun: corona – Sun: flares – Sun: transition region – Sun: UV radiation – Sun: X-rays, gamma rays

## 1. Introduction

In the widely accepted “standard” solar flare model, pre-flare energy is stored in non-potential magnetic fields. This energy is released by magnetic reconnection, during which nonthermal particle beams are accelerated and the plasma near the reconnection site (typically thought to be in the corona) is heated. The released energy is subsequently transported to and deposited in the chromosphere either by nonthermal beam particles or by thermal conduction from the directly-heated coronal source. When the flux of energy into the chromosphere exceeds what can be shed by radiative losses, the chromosphere heats and expands. Because the density in the overlying corona is less than that in the chromosphere and photosphere, expansion of the heated chromospheric material produces larger upward than downward velocities. This process, now known as “chromospheric evaporation” (see Antonucci *et al.* 1999 and Bornmann 1999 for reviews) was first described by Neupert (1968). Upward velocities around 100 km s<sup>-1</sup> or more in emission lines formed at temperatures  $T \gtrsim 10$  MK during flares have long been interpreted as evidence for this phenomenon (Antonucci *et al.* 1982; Antonucci & Dennis 1983; Zarro *et al.* 1988; Fludra *et al.* 1989; Canfield *et al.* 1990; Doschek 1990; Mariska *et al.* 1993; Mariska 1994;

Silva *et al.* 1997; Brosius 2003, 2009; Brosius & Phillips 2004; Milligan & Dennis 2009). The evaporated material fills magnetic loops with hot, dense, plasma that eventually cools and fades as the flare decays.

Investigations of solar flares based on rapid cadence (7.2 to 9.8 s) stare spectra obtained with the Coronal Diagnostic Spectrometer (CDS; Harrison *et al.* 1995) aboard the *Solar and Heliospheric Observatory (SOHO)* spacecraft showed that intensities of emission lines formed in plasma at transition region temperatures (*e.g.*, the O V line at 629.7 Å, formed at  $T \approx 0.25$  MK) began to increase several minutes earlier than intensities of emission lines formed in plasma at flare temperatures (*e.g.*, the Fe XIX line at 592.2 Å,  $T \approx 8$  MK; Brosius & Phillips 2004; Brosius 2009; Brosius & Holman 2009, 2010). Statistically significant relative Doppler velocities also typically accompany the intensity increases in the emission lines. This time difference is consistent with the chromosphere being heated predominantly by beams of nonthermal particles; if the chromosphere were heated predominantly by thermal conduction from a directly-heated  $T \sim 10$  MK source, that source would have appeared first, before the chromosphere had a chance to respond to thermal conduction from it. When possible, CDS stare spectra were combined with X-ray observations from the *Ramaty High Energy*

*Solar Spectroscopic Imager (RHESSI; Lin et al. 2002; Hurford et al. 2002; Smith et al. 2002)*. *RHESSI* combines full disk solar imaging, wide energy coverage (3 to 17,000 keV), and high sensitivity for investigating both nonthermal and high temperature ( $\gtrsim 10$  MK) thermal emission from flares and microflares.

A shortcoming of the rapid cadence CDS stare studies is their comparatively limited field of view. In previous investigations, this was partly compensated by acquiring full disk 195 Å images at 12 minute cadence from the Extreme-ultraviolet Imaging Telescope (EIT; Delaboudinière et al. 1995) aboard *SOHO*. Nevertheless, there remained some uncertainty as to what was going on outside the CDS field of view. This spatial uncertainty can be eliminated by using the Atmospheric Imaging Assembly (AIA; Lemen et al. 2011) aboard the *Solar Dynamics Observatory (SDO)* satellite: since AIA obtains snapshots of the entire Sun in each of its EUV channels every 12 s, the start of any given flare or microflare is never hidden outside AIA's field of view, and can always be determined for each channel within the instrumental cadence of 12 s.

Launched on 2010 February 11, *SDO* observes the Sun nearly continuously in an inclined geosynchronous orbit from which its extremely high data rate is enabled through contact with a single dedicated ground station. AIA contains seven EUV channels roughly centered on the following wavelengths, designed to include solar radiation in and around strong emission lines from the following ions: 304 Å (includes emission from He II, formed at  $T \approx 0.05$  MK), 171 Å (Fe IX at  $T \approx 0.63$  MK), 211 Å (Fe XIV at  $T \approx 1.9$  MK), 335 Å (Fe XVI at  $T \approx 2.7$  MK), 94 Å (Fe XVIII at  $T \approx 6.3$  MK), 193 Å (Fe XII at  $T \approx 1.4$  MK and Fe XXIV at  $T \approx 19$  MK), and 131 Å (Fe VIII at  $T \approx 0.37$  MK, Fe XX at  $T \approx 9.1$  MK, Fe XXI at  $T \approx 10$  MK, and Fe XXIII at  $T \approx 14$  MK). Wavelength response functions for each channel were measured in the laboratory prior to *SDO*'s launch (Boerner et al. 2011), and are available through SSWIDL. Temperature response functions for each channel were derived (Boerner et al. 2011) by combining the measured wavelength response functions with emission line data from the CHIANTI database (Dere et al. 1997, 2009), and are also available through SSWIDL. Based on AIA's powerful combination of continuous, rapid cadence (12 s), full disk EUV imaging and wide range of temperature sensitivity, its observations are expected to provide new information on solar activity including the temperature dependence of flare start times. The use of timing differences among various EUV channels to assess the roles of thermal conduction and nonthermal particle beams in transporting the energy that drives chromospheric evaporation is one of AIA's objectives (Concept Study Report; see <http://aia.lmsal.com/public/CSR.htm>).

Here we present coordinated AIA and *RHESSI* observations of a *GOES* B4.8 microflare on 2010 July 31. AIA light curves averaged over selected areas within the active region are expected to reveal whether timing differences exist among the various channels and, hence, the importance of thermal conduction relative to nonthermal particles in transporting the energy that drives chromospheric evaporation. To help with the interpretation of these and additional AIA flare observations, we characterize the expected response of AIA's EUV channels to the onset of solar flare brightenings by combining (1) AIA's temperature response functions (Boerner et al. 2011) with (2) existing EUV spectral line flare data obtained with CDS on timescales comparable to AIA's image cadence (Brosius & Phillips 2004). In §2 we describe the 2010 July 31 flare observational results, in §3 we present the expected temporal response of AIA's EUV channels

based on CDS flare spectra **obtained on 2001 April 24**, in §4 we discuss implications of this work, and in §5 we summarize our findings.

## 2. Observation and Analysis of Microflare on 2010 July 31

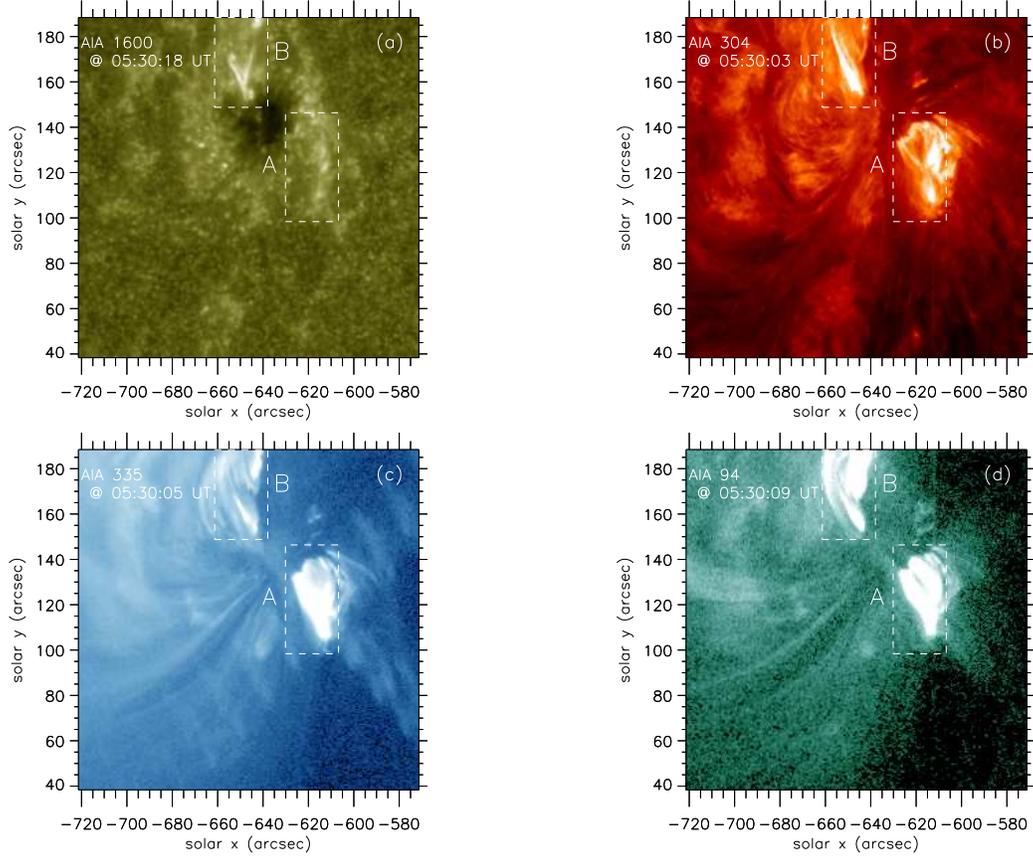
We observed NOAA Active Region 11092 on 2010 July 31 with AIA from 05:00 to 05:45 UT and with *RHESSI* from 05:00 to 05:40 UT, during which the *GOES* satellite observed a B4.8 microflare from 05:17 to 05:36 UT. We used AIA's online "cutout service" (which provides Level 1.5 data; see Lemen et al. 2011) to extract high spatial resolution ( $0''.6$  pixels), rapid cadence (12 s), coaligned  $2'.5 \times 2'.5$  images centered on  $(-646''.3, +113''.4)$  from all seven of AIA's EUV channels as well as its 1600 Å channel (at 24 s cadence).

Figure 1 displays a sample of  $2'.5 \times 2'.5$  images of AR 11092 in four different channels observed by AIA on 2010 July 31. The source areas A and B in which two successive microflare bursts originate are separated from each other by  $30''$  or more but both connected to the sunspot toward the center of the field of view. Movies in all eight of the AIA channels (two of which are available online) indicate that the microflare started within the area outlined by the lower dashed rectangle (A), and occurred later within the area outlined by the upper dashed rectangle (B). It is possible that the B4.8 microflare observed by *GOES* was actually a pair of closely spaced, closely timed sympathetic microflares.

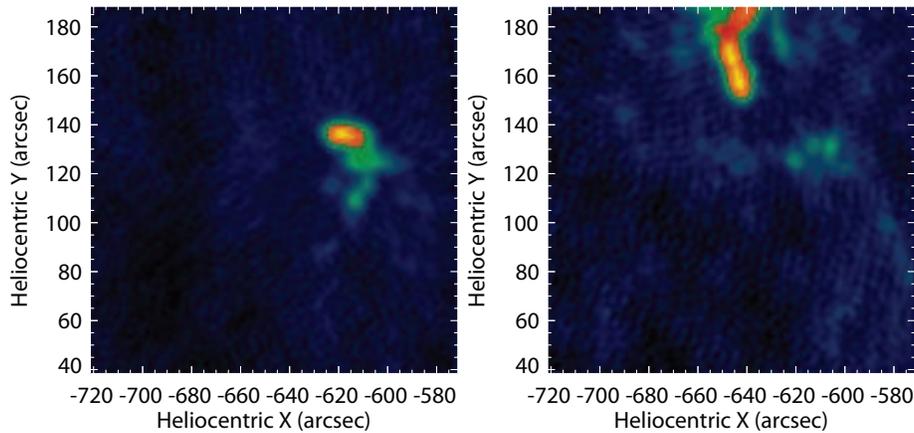
Figure 2 displays a pair of *RHESSI* 3-25 keV images that show the two spatially separated sources of brightenings that are seen in the AIA images of the microflare. The time interval for each image is indicated in the figure captions. We integrated over 1-minute intervals because the count rates were low. These images were obtained using the CLEAN technique and *RHESSI* front detectors 1 - 9, giving a spatial resolution  $\sim 3''$ .

Figures 3 and 4 both show 3-12 and 12-25 keV *RHESSI* light curves. Unfortunately, a particle event in Earth's magnetosphere contaminated these with gradual humps that peak toward the center of *RHESSI*'s observing window, so while the resulting increased background is not expected to significantly alter the start times or peak times of the flare bursts, it effectively hides any early gradual increase to which *RHESSI* may have been sensitive. Note the mound in the 12-25 keV emission atop which bursts starting at 05:17 and 05:27 UT are seen in Figures 3 and 4. *GOES*, however, is unaffected by this, and shows a flat background until the start of its first recorded burst at 05:17:32 UT.

Figure 3 displays light curves for all seven of AIA's EUV wavelength channels, plus that of 1600 Å, averaged over area A in Figure 1 to ensure that the microflare's starting position is included in the light curves. Solid vertical lines indicate the start of an increase in EUV emission observed by AIA no later than 05:11:34 UT, a peak in the early (precursor?) brightening at 05:13:30 UT, the start of the first impulsive rise of AIA emission at 05:16:44 UT, and the start of the first burst recorded by *GOES* at 05:17:32 UT. The first impulsive peak recorded by AIA occurred at 05:18:30 UT (simultaneously for all wavebands, within the uncertainty of the cadence), and the rise toward a secondary peak began at 05:19:34 UT. Figure 4 displays light curves for the same AIA wavelength channels averaged over area B in Figure 1. The *GOES* and *RHESSI* light curves from the start of the microflare component observed in area B are not as pristine as those from the start of the microflare component observed in area A because they are superimposed on declining emission from A,



**Fig. 1.**  $2'.5 \times 2'.5$  AIA images of AR 11092 in (a) the 1600 Å channel at 05:30:18 UT, (b) the 304 Å channel at 05:30:03 UT, (c) the 335 Å channel at 05:30:05 UT, and (d) the 94 Å channel at 05:30:09 UT on 2010 July 31. The dashed rectangles labeled “A” and “B” outline areas within which successive bursts were observed with AIA (first in A, later in B) and *RHESSI* during the *GOES* B4.8 microflare, and for which we derived the spatially averaged AIA light curves shown in Figures 3 and 4. Temporal evolution in the 304 and 94 Å channels can be seen in the movies available online.

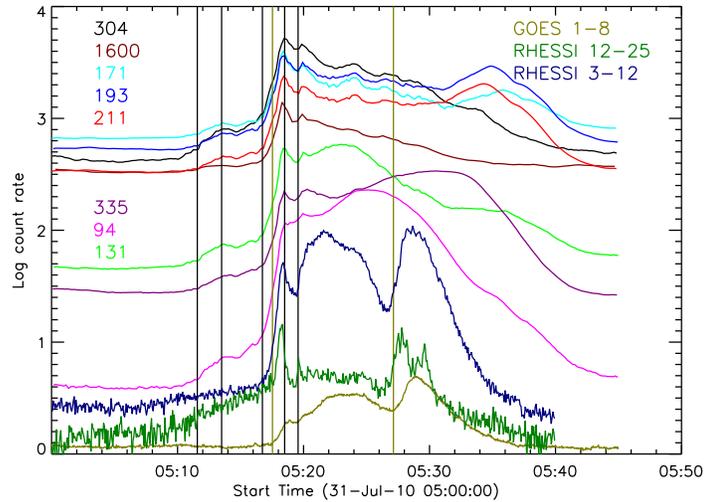


**Fig. 2.**  $2'.5 \times 2'.5$  *RHESSI* 3–25 keV images of the *GOES* B4.8 microflare integrated over 1-minute time intervals starting at 05:17:40 UT (left panel), and at 05:27:20 UT (right panel) on 2010 July 31.

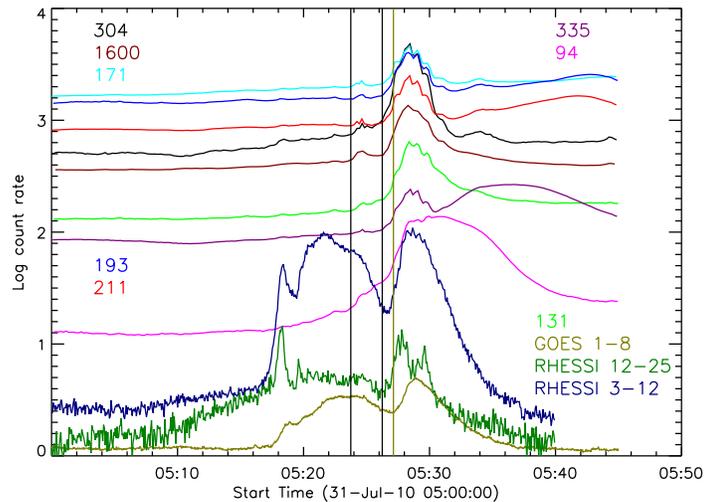
but they show that the rise in the AIA light curves at that location begins before the start of the second burst in the *GOES* or *RHESSI* light curves. Solid vertical lines indicate the start of an increase in EUV emission no later than 05:23:45 UT in B, the start of a rapid rise in EUV emission no later than 05:26:15 UT, and the start of the second burst recorded by *GOES* at 05:27:08 UT.

Eventually we see different temporal behaviors in the different channels. For example, post-impulsive maxima occur at pro-

gressively later times in channels that were designed to observe bright lines of successively cooler ions, consistent with expectations of a cooling plasma after the cessation of energy deposition. During this post-impulsive cooling phase, the AIA channel that observes Fe XXI flare emission (131 Å) peaks first (a somewhat broad peak centered around 05:23 UT), followed by 94 Å (05:25 UT), and so on down to 171 Å (05:36 UT). The fact that the 193 Å channel does not show a similar post-impulsive peak



**Fig. 3.** Light curves (log of the count rate for AIA, arbitrary units for *GOES*, log of the count rate for *RHESSI*, but shifted downward by 2.2 for the 3-12 keV light curve, and 1.65 for the 12-25 keV light curve) from eight wavelength channels on AIA averaged over area A (see Figure 1) in which the brightening first occurred during the B4.8 microflare on July 31. Solid vertical lines indicate the start of an increase in EUV emission observed by AIA no later than 05:11:34 UT, a peak in the early (precursor?) brightening at 05:13:30 UT, the start of an impulsive rise in AIA emission at 05:16:44 UT, the start of the first burst recorded by *GOES* at 05:17:32 UT, the first impulsive peak detected by AIA at 05:18:30 UT (simultaneously for all wavebands, within the uncertainty of the cadence), the rise toward a secondary peak at 05:19:34 UT, and the start of the second burst recorded by *GOES* at 05:27:08 UT.



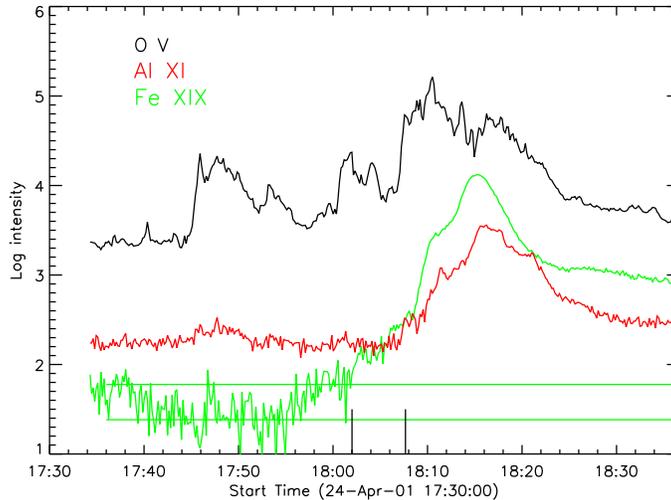
**Fig. 4.** Light curves (log of the count rate for AIA, arbitrary units for *GOES*, log of the count rate for *RHESSI*, but shifted downward by 2.2 for the 3-12 keV light curve, and 1.65 for the 12-25 keV light curve) from eight wavelength channels on AIA averaged over area B (see Figure 1) in which the brightening started later during the B4.8 microflare on July 31. Solid vertical lines indicate the start of an increase in EUV emission observed by AIA no later than 05:23:45 UT, the start of a rapid rise in EUV emission no later than 05:26:15 UT, and the start of the second burst recorded by *GOES* at 05:27:08 UT.

prior to that of the 131 Å emission indicates that the maximum flare temperature remained smaller than that required to produce Fe XXIV.

### 3. Expected Temporal Response of AIA Based on CDS Flare Spectra from 2001 April 24

To quantify the rapid cadence (9.8 s) temporal behavior of the solar flare atmosphere over a wide range of temperature ( $4.67 \leq \log T \leq 6.89$ ), we use the integrated intensities ( $\text{ergs cm}^{-2} \text{s}^{-1} \text{sr}^{-1}$ ) derived by Brosius & Phillips (2004) for the following nine emission lines measured at the commencement site (an area in

and around one of the flare loop's footpoints) of a *GOES* M2.3 flare in Active Region 9433 on 2001 April 24: He II 303.78 Å ( $\log T = 4.67$ ), O III 599.60 Å ( $\log T = 4.96$ ), O IV 554 Å (a self-blend of four lines at 553.33, 554.08, 554.51, and 555.26 Å, formed at  $\log T = 5.24$ ), O V 629.76 Å ( $\log T = 5.37$ ), Ne VI 558.59 Å ( $\log T = 5.63$ ), Ca X 557.76 Å ( $\log T = 5.87$ ), Mg X 609.79 Å ( $\log T = 6.05$ , from which its O IV blend was removed based on the density-insensitive intensity ratio  $\text{O IV } 609.83 / (553.33 + 554.08 + 554.51 + 555.26) = 0.1645 \pm 0.0022$  derived with CHIANTI v. 6.0.1; Dere *et al.* 1997, 2009), Al XI 550.04 Å ( $\log T = 6.17$ ), and Fe XIX 592.23 Å ( $\log T = 6.89$ ).



**Fig. 5.** Light curves of O V 629.7 Å, Al XI 550.0 Å, and Fe XIX 592.2 Å obtained with *SOHO*'s CDS in rapid cadence (9.8 s) stare mode during the 2001 April 24 *GOES* M2.3 flare reported by Brosius & Phillips (2004). We use these and light curves from six additional spectral lines to measure average preflare intensities and flare intensity enhancement factors as described in the text. Solid horizontal lines indicate  $\pm 1\sigma$  from the average pre-flare Fe XIX noise. The long vertical tick mark at 18:02:00 UT indicates the time at which the Fe XIX intensity exceeds and remains continuously greater than its pre-flare noise level by more than  $1\sigma$ , and the long vertical tick mark at 18:07:40 UT indicates the time at which *Yohkoh* hard X-ray count rates became statistically significant (with a peak at 18:10:40 UT).

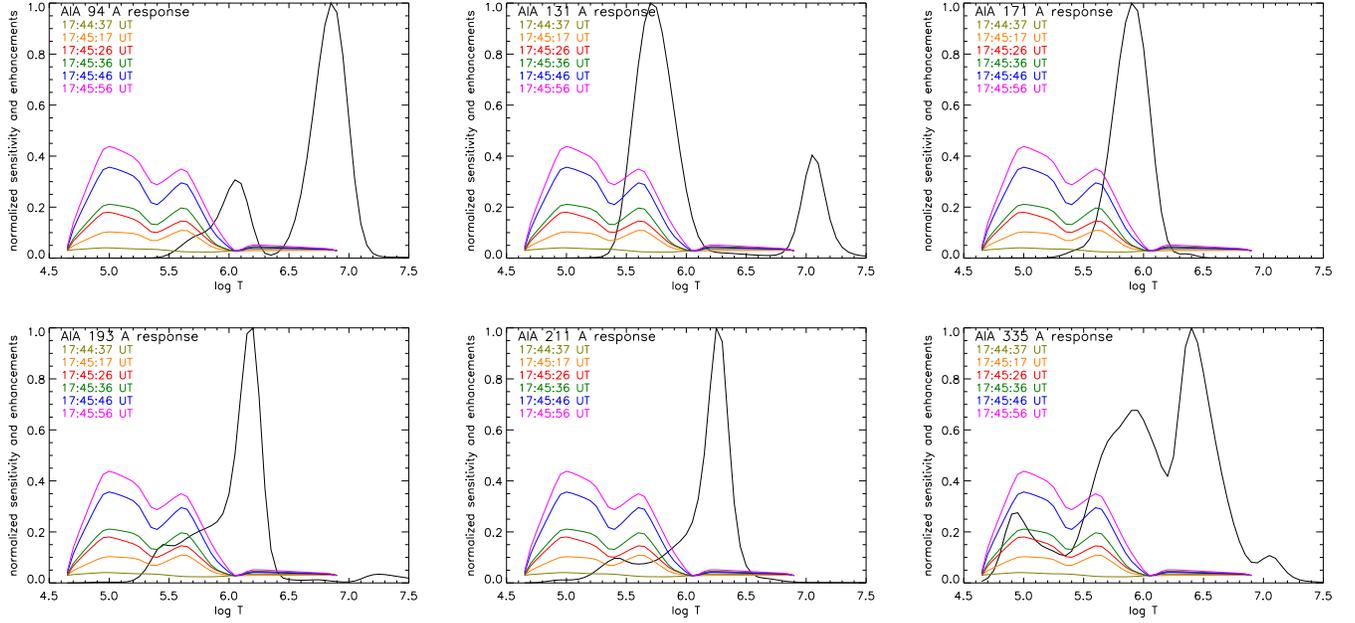
See Figure 5. Using the 51 exposures whose start times occurred between 17:36:05 and 17:44:18 UT, we derived pre-flare average intensities and their corresponding 1-sigma scatters. All of the lines except that of Fe XIX at 592.2 Å were observed during this pre-flare interval; for Fe XIX the pre-flare average intensity is defined by the noise. We consider the detection of Fe XIX emission to be reliable only when the intensity derived from fits to the line profile exceeds and continuously remains greater than the pre-flare average noise level by more than  $1\sigma$  (which occurs at 18:02:00 UT; see Figure 5). Upward velocities approaching  $-40 \text{ km s}^{-1}$  in only the transition region lines of O III, O IV, O V, and He II during the “precursors” (17:44:30 - 17:57:10 and 18:00:30 - 18:05:30 UT) indicated gentle chromospheric evaporation; downward velocities of  $+40 \text{ km s}^{-1}$  in those same lines during the impulsive phase, when Fe XIX showed maximum upward velocities of  $-64 \text{ km s}^{-1}$ , indicated explosive evaporation.

For each exposure beginning with the start of the first precursor, we calculate an intensity enhancement factor for each emission line as the ratio of the line’s intensity in the given exposure to its pre-flare average value. Since the Fe XIX line emission is considered to be noise until its intensity continuously exceeds the pre-flare average noise by more than  $1\sigma$ , we define its enhancement factor as 1 until that time. Thus for each 9.8 s interval we calculate 9 enhancement factors, and we assume that each corresponds to its respective emission line’s formation temperature (listed above). We use a spline fit to place these onto the finer temperature grid (0.05 dex) of the AIA temperature response functions (Boerner *et al.* 2011). In this way we obtain enhancement factors as functions of temperature for  $4.65 \leq \log T \leq 6.90$ . See Figure 6, which shows normalized AIA temperature response functions for all of the EUV channels except 304 Å, along with normalized enhancement factor curves for six CDS exposures up to and including the first peak during the flare’s first precursor. In Figure 6 and subsequent figures, the AIA temperature response functions are normalized to the maximum value within their respective channels. The EUV spectral line enhancement factors are normalized to the largest overall

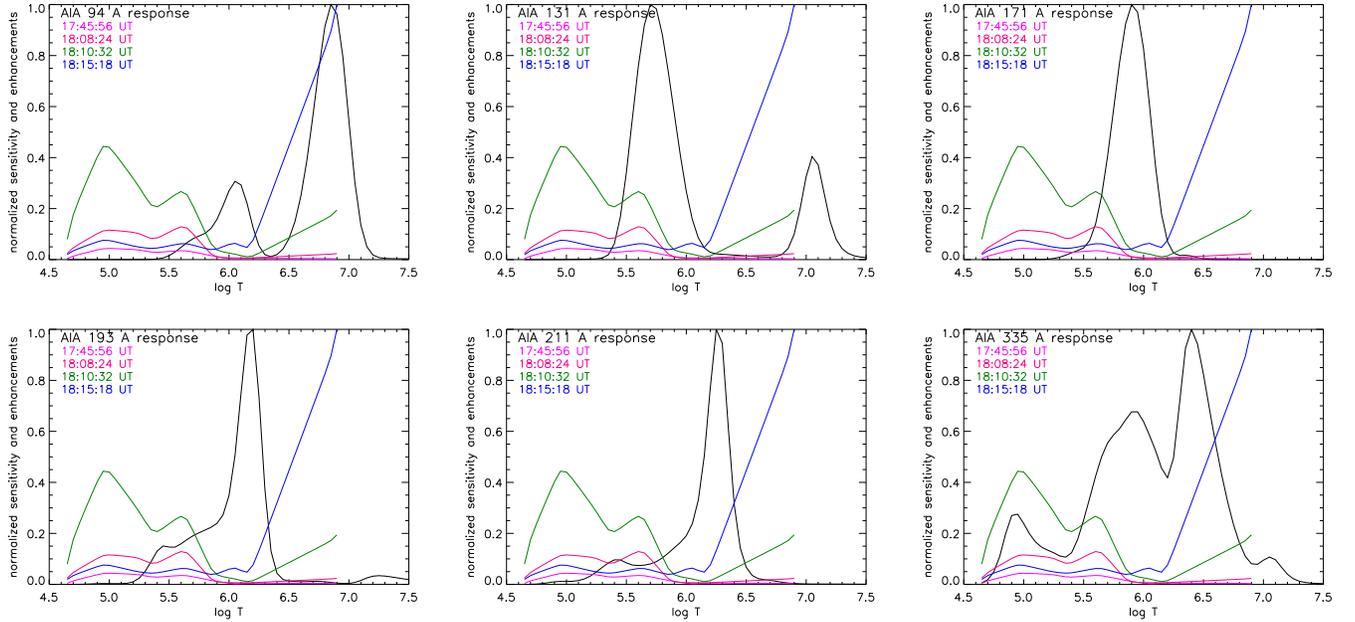
value during the flare (325, which occurs for Fe XIX during its peak at 18:15:18 UT), multiplied by 10 in Figure 6 to improve visibility. The point of Figure 6 is that the majority of the intensity enhancements observed early during the first precursor occur at temperatures less than about  $\log T = 5.8$  (0.6 MK), with the largest values occurring around  $\log T = 5.0$  (0.1 MK).

Figure 7 is similar to Figure 6, and shows normalized enhancement factors (in this case *not* multiplied by a factor of 10) for four CDS exposures during the flare. These include the first peak during the first precursor (17:45:56 UT, which is also shown in Figure 6), the start of the impulsive increase of the Fe XIX emission at 18:08:24 UT, the peak intensity of the cool, transition region emission lines at 18:10:32 UT, and the peak intensity of the Fe XIX line observed by CDS at 18:15:18 UT (the maximum to which all of the CDS enhancements are normalized). It is evident that although intensity enhancements due to low temperature emission dominate early during the flare, by the time of the maximum O V intensity the contributions from hot flare emission (approaching 10 MK) are significant while at the time of the maximum Fe XIX intensity they dominate the low temperature contributions.

Figure 8 shows fractional contributions to the total brightness increase in AIA’s 94, 131, 171, 193, 211, and 335 Å channels due to plasma with  $\log T \leq 5.85$ . Because the temperature range for this calculation is limited to that available to CDS ( $4.67 \leq \log T \leq 6.89$ ), we define the fractional contribution for each frame in this figure as the ratio of the enhancement summed over the interval from  $\log T = 4.65$  to 5.85 to that summed over the interval from  $\log T = 4.65$  to 6.90. The curves show that contributions to the total brightness increase in each channel are greatest for low-temperature emission until about 18:10 UT, when contributions from emission formed at larger temperatures begin to dominate. The lower, darker curve in Figure 8a was calculated with the 94 Å channel’s nominal temperature response function (Boerner *et al.* 2011) obtained through SSWIDL, while the upper, lighter curve was calculated with a “corrected” version of



**Fig. 6.** Nominal, normalized AIA temperature response functions (Boerner *et al.* 2011), with CDS enhancement factors (normalized to the largest value which occurs for Fe XIX, then multiplied by 10 to improve visibility in these frames) for six of the first ten CDS exposures, beginning with the first exposure during the first precursor and ending with the first peak during the first precursor of the M2.3 flare. The interval between exposures is 9.8 s. The majority of the intensity enhancements occur at temperatures less than  $\log T = 5.8$ , with the largest values at  $\log T = 5.0$  (for the last four CDS exposures in this series). The greatest contributions to AIA’s brightenings occur where the product of the intensity enhancements and the response functions are largest.

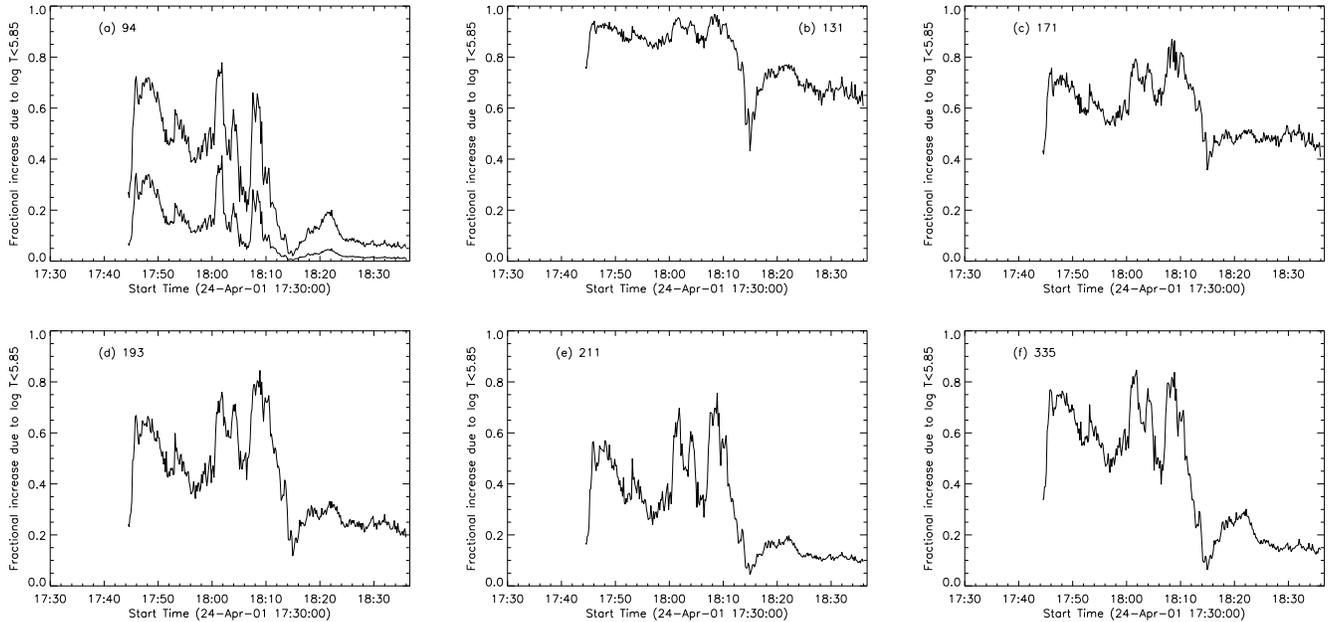


**Fig. 7.** Nominal, normalized AIA temperature response functions, with CDS enhancement factors (normalized to the largest value which occurs for Fe XIX) for four times during the M2.3 flare, including (1) the first peak of the first precursor at 17:45:56 UT (the top-most CDS curve in Figure 6, but here not multiplied by 10), (2) the start of the impulsive increase of the Fe XIX emission at 18:08:24 UT, (3) the peak intensity of the cool, transition region emission lines at 18:10:32 UT, and (4) the peak intensity of the Fe XIX line observed by CDS at 18:15:18 UT (the maximum to which all of the CDS enhancements are normalized). The interval between exposures is 9.8 s.

this function in which the temperature response was multiplied by 5 for  $\log T \leq 5.85$  (see Discussion).

Based on our combination of the nominal AIA temperature response curves available through SSWIDL (Boerner *et al.* 2011) with the (temperature-dependent) intensity enhancement

factors derived from the CDS spectral line data of the M2.3 flare on 2001 April 24 (Brosius & Phillips 2004), we estimate that AIA would have observed count rate increases in its 94 Å channel by factors of 1.87 at the first peak of the flare’s first precursor, 6.94 by the start of the impulsive rise of the Fe XIX emission,



**Fig. 8.** Fractional contribution to total brightness increase in six AIA EUV channels attributed to plasma with  $\log T \leq 5.85$ . For this investigation our temperature range is constrained to that available to CDS, namely,  $4.67 \leq \log T \leq 6.89$ . We increased AIA’s 94 Å channel’s low-temperature ( $4.65 \leq \log T \leq 5.85$ ) sensitivity by a factor of 5 to compensate for its known *underestimate* of its low-temperature sensitivity; this is plotted as the upper (lighter color) curve in frame (a), as described further in the text.

40.6 at the overall maximum in the “cool” (*e.g.*, O V) line emission, and 186.5 at the time of the overall maximum in the Fe XIX flare line emission. These and corresponding values at the same times in the 131, 171, 193, 211, and 335 Å channels are listed in Table 1. By way of comparison, AIA channel enhancement factors observed on 2010 July 31 at the peak (05:13:30 UT) of the earliest flare (precursor?) brightening and the first impulsive peak (05:18:30 UT) of the light curves from area A (see Figure 1) are given in Table 1.

All of AIA’s EUV channels are sensitive to emission around  $\log T \approx 5.5$  (see Figures 6 and 7). This means that any flare (like that observed by Brosius & Phillips 2004) in which emission from plasma at  $\log T \approx 5.5$  brightens significantly earlier (by more than AIA’s 12 s cadence) than emission from plasma at “flare” temperatures ( $\log T \approx 7$ ) will brighten simultaneously in AIA’s 94, 131, 171, 193, 211, and 335 (and 304) Å channels. As can be seen from AIA’s temperature response functions shown in Figures 6 and 7, all of those channels are also sensitive to emission from plasma at  $\log T \approx 6$ . Thus, while a simultaneous brightening in all of AIA’s EUV channels would be consistent with previous spectroscopic observations in which emission from plasma at  $\log T \approx 5.5$  is the first to brighten during a flare (*e.g.*, Brosius & Phillips 2004; Brosius 2009; Brosius & Holman 2009, 2010), the possibility would remain that the actual temperature of the emitting plasma is close to  $\log T \approx 6$  or perhaps spread over a range from  $\log T \approx 5$  to  $\log T \approx 6$ . Thus for AIA flare observations in which all of AIA’s EUV channels brighten simultaneously we encounter some ambiguity regarding the temperature distribution of the emitting plasma. Nevertheless, a simultaneous brightening in all of AIA’s EUV channels is consistent with chromospheric heating and evaporation being driven by beams of nonthermal particles.

If, on the other hand, emission from plasma at  $\log T \approx 7$  (the result of direct coronal heating at the reconnection site) brightened significantly before emission from plasma at  $\log T \approx 5.5$ ,

then only the AIA channels that are sensitive to emission from such hot plasma (94 and 131 Å) will brighten significantly while others (193 and 335 Å) may show a smaller brightening and the remainder (171 and 211 Å) will show no effect. This enables AIA to distinguish which mechanism (thermal conduction or nonthermal particle beam) dominates the transport of flare energy from its release site (typically thought to be in the corona) to the chromosphere. By way of quantifying AIA’s flare response for a case in which reconnection-heated, high-temperature (10 MK) plasma appears before the brightening of plasma at lower temperatures, we consider a synthetic case in which only the intensity of the hot flare plasma ( $6.8 \leq \log T \leq 6.9$ ) is enhanced early during the flare. Based on calculations similar to those described above for the combination of AIA temperature response curves and CDS flare spectra, we find that the 94 Å channel would brighten by a factor of 4.51 when the intensity enhancement factor is 10, and by a factor of 12.3 when it is 30. Similarly, the 131 Å channel would brighten by factors of 1.18 and 1.57 for the same intensity enhancement factors. For the 171 Å channel we get 1.00 and 1.00, for the 193 Å channel we get 1.03 and 1.09, for the 211 Å channel we get 1.02 and 1.08, and for the 335 Å channel we get 1.18 and 1.57.

## 4. Discussion

Figure 3 shows that all of the AIA light curves from area A in Figure 1 started to rise at the same time, nearly 6 minutes before *RHESSI* or *GOES* detected any increase due to the microflare, while Figure 4 shows that all of the AIA light curves from area B started their rise more than 3 minutes before *RHESSI* or *GOES* detected an increase. This is consistent with expectations presented in §3, in which the simultaneous increase observed in AIA’s 94, 131, 171, 193, 211, and 335 Å channels is due to emission from plasma at temperatures between 0.1 and 0.7 MK. This

is consistent with nonthermal electrons starting to heat the chromosphere nearly 6 minutes before *RHESSI* or *GOES* detected microflare emission. If the chromosphere were heated by thermal conduction, an earlier increase in high-temperature ( $\sim 10$  MK) emission is expected to have been seen by *RHESSI*, *GOES*, or the AIA channels that were designed with the greatest sensitivity to emission from high-temperature flare plasma (Fe XVIII at 6.3 MK in the 94 Å channel, Fe XXI at 10 MK in the 131 Å channel, Fe XXIV at 19 MK in the 193 Å channel, and Ni XXII at 10 MK and Fe XXI at 10 MK in the 335 Å channel). Since this was not the case, we find no observational evidence to suggest that thermal conduction heated the chromosphere or drove evaporation at the start of the microflare (in either area A or area B).

The *RHESSI* data show no X-ray emission before about 15:16 UT. After this time, the 12 - 25 keV light curve (see Figure 3) displays the impulsive character of bremsstrahlung from non-thermal electrons. The first X-ray peak, beginning between 5:17 UT and 5:18 UT and ending between 5:19 UT and 5:20 UT, may in fact contain significant emission from nonthermal electrons down to the lowest observed photon energy of 3 keV. Spectral analysis to investigate the likelihood that the higher energy photons (and a significant fraction of the lower energy X-ray photons in the first peak) were from nonthermal electrons was inconclusive because of the low level of emission above 12 keV and the uncertainty in the background subtraction. Despite *RHESSI*'s observation of no hard X-ray emission before 5:16 UT, nonthermal electrons could still have provided the power input required to heat the chromospheric plasma at these earlier times (e.g., Siarkowski et al. 2009; Falewicz et al. 2011).

In addition to the light curves from areas A and B, we also obtained light curves averaged over the entire  $2'.5 \times 2'.5$  field of view that we had selected from the AIA cutout service. Because these light curves are averaged over large areas that do not participate in the flare brightenings, features in them appear relatively weakened or suppressed when compared with their counterparts from areas A and B. Although this makes the starting times of some of the features indicated in Figures 3 and 4 more difficult to determine, the relative timings remain unchanged. Further, and perhaps most important, is the fact that the light curves averaged over this large area yield no evidence of a “hot” ( $T \sim 10$  MK) start that may have occurred outside areas A and B above. This provides further support for our conclusion that flare energy was transported by nonthermal particle beams rather than by thermal conduction during the 2010 July 31 B4.8 microflare.

In a recent theoretical investigation O’Dwyer *et al.* (2010) computed the contributions of spectral lines and continuum emission to AIA’s EUV channels for different regions of the solar atmosphere (coronal hole, quiet Sun, active region, and flare plasma), and concluded that no waveband is strictly isothermal: emission from plasma at a variety of temperatures contributes significantly to each. Although Table 1 of O’Dwyer *et al.* (2010) indicates that AIA’s 94, 131, and 193 Å channels are dominated by Fe XVIII, Fe XXI, and Fe XXIV, respectively, during a flare, those calculations were performed with a differential emission measure obtained during the *decay* of an M2 flare (Dere & Cook 1979). This DEM did not include the flare impulsive phase, when bursts of emission from plasma at transition region temperatures (0.1 - 0.5 MK) are particularly bright (see, e.g., Fisher *et al.* 1985a,b,c and CDS light curves presented by Brosius 2003, Brosius & Phillips 2004, and Brosius & Holman 2007). Thus, as we have demonstrated in §3 above, emission from transition region plasma will dominate in all of AIA’s EUV channels during

a flare’s precursors and early impulsive rise, while emission from plasma at temperatures around 10 MK will dominate later during the event. In an even more recent comparison of AIA images of a non-flaring loop footpoint and leg with coordinated EUV spectra from *Hinode*’s Extreme-ultraviolet Imaging Spectrometer (EIS, Culhane *et al.* 2007), Del Zanna *et al.* (2011) showed that emission from lines formed at transition region temperatures can be significant for all of AIA’s EUV channels, and dominant in some cases. Their observation of numerous unidentified lines in the EIS spectra underscored the need for further work on the relevant atomic data before AIA images can be reliably used for plasma diagnostic purposes.

A number of emission lines that are found, for example, in the Kelly (1987) Finding List at wavelengths between 93.5 and 94.5 Å are absent from the CHIANTI v.6.0.1 database (Dere *et al.* 1997, 2009) on which the AIA temperature response functions as of this writing are based (Boerner *et al.* 2011). These include 11 lines of Al V (log  $T = 5.42$ ), 1 line of O VI (log  $T = 5.46$ ), 1 line of Na VI (log  $T = 5.65$ ), 2 lines of Al VII (log  $T = 5.79$ ), 10 lines of Na VII (log  $T = 5.78$ ), and 2 lines of Na VIII (log  $T = 5.86$ ). Any or all of these might significantly increase the low-temperature sensitivity of the 94 Å channel’s temperature response function. For demonstration purposes, we multiplied the existing, nominal 94 Å channel’s temperature response function (Boerner *et al.* 2011; available through SSWIDL, and used to derive the lower, darker curve in Figure 8a) by a factor of 5 for log  $T \leq 5.85$ . This value is at the low end of the  $6.7 \pm 1.7$  range in correction factor proposed by Aschwanden & Boerner (2011), but it is difficult to say what the correction factor really is without the additional relevant atomic physics data that need to be generated and included in the CHIANTI database. We repeated the calculation of the fractional contribution to the total brightness increase in AIA’s 94 Å channel due to plasma with log  $T \leq 5.85$ , and display the result as the upper, lighter-color curve in Figure 8a. Note that the greater sensitivity to low temperature emission translates to a larger fractional contribution to flare brightening due to this emission. It does not, however, alter our conclusions regarding the simultaneity of flare brightenings in AIA’s EUV channels. Further, it appears likely that the 131 Å channel’s temperature response function will undergo a similar revision once the 7 lines of O V (log  $T = 5.37$ ), 4 lines of Al V (log  $T = 5.42$ ), 8 lines of Na V (log  $T = 5.46$ ), and 8 lines of Ne V (log  $T = 5.47$ ) in the Finding List between wavelengths of about 128 and 132 Å are also available through CHIANTI to be included in the 131 Å channel’s temperature response function.

Further studies that compare AIA images with coordinated CDS, EIS, or EUNIS (Extreme Ultraviolet Normal Incidence Spectrograph; Brosius *et al.* 2008, Wang *et al.* 2010, 2011) spectral observations that cover AIA channels accessible to those instruments will help to better quantify the temperature responses of those channels by directly showing all of the spectral lines that produce the observed count rates in selected areas within a variety of solar features.

## 5. Summary

Coordinated AIA and *RHESSI* observations of a *GOES* B4.8 microflare on 2010 July 31 show that emission in all seven of AIA’s EUV channels brightened simultaneously nearly 6 minutes before *RHESSI* or *GOES* detected emission from plasma at temperatures around 10 MK. To help interpret these and AIA flare observations in general, we characterized the expected temporal responses of AIA’s 94, 131, 171, 193, 211, and 335 Å channels to

**Table 1.** AIA Intensity Enhancement Factors at Selected Times During Flares.

Channel	Based on AIA and CDS <sup>a</sup>				AIA Observed <sup>b</sup>	
	pre peak <sup>c</sup>	rise XIX <sup>d</sup>	O V max <sup>e</sup>	max XIX <sup>f</sup>	1st peak <sup>g</sup>	imp. peak <sup>h</sup>
94 Å	1.87	6.94	40.6	186.5	1.85	29.2
131 Å	7.51	22.9	51.2	25.0	1.63	11.5
171 Å	4.65	10.7	25.8	17.3	1.33	6.00
193 Å	3.58	9.14	22.2	27.4	1.38	6.74
211 Å	3.08	7.45	21.5	45.2	1.33	7.10
335 Å	4.62	12.7	40.8	67.9	1.35	7.85

**Notes.** <sup>(a)</sup> Expected AIA intensity enhancement factor based on nominal temperature response functions (Boerner *et al.* 2011) combined with intensity enhancement factors measured with CDS spectra during the M2.3 flare on 2001 April 24 (Brosius & Phillips 2004). <sup>(b)</sup> Values observed by AIA in area A of Figure 1 during the microflare on 2010 July 31. <sup>(c)</sup> At the first peak of the flare’s first precursor (17:45:56 UT). <sup>(d)</sup> At the impulsive rise of the Fe XIX emission (18:08:24 UT). <sup>(e)</sup> At the maximum intensity observed in the cool (*e.g.*, O V) emission lines (18:10:32 UT). <sup>(f)</sup> At the maximum of the Fe XIX emission (18:15:18 UT). <sup>(g)</sup> At the first peak after the microflare’s onset (05:13:30 UT). <sup>(h)</sup> At the microflare’s first impulsive peak (05:18:30 UT).

solar flare brightenings by combining (1) AIA’s nominal temperature response functions available through SSWIDL (Boerner *et al.* 2011) with (2) EUV spectral line flare data obtained with CDS on 2001 April 24 (Brosius & Phillips 2004) on timescales comparable to AIA’s image cadence. The line brightenings that were observed early during the CDS flare occurred at temperatures less than about 0.7 MK, with the largest values around 0.1 MK; this is consistent with the flare’s energy transport being dominated by nonthermal particle beams. Because all of AIA’s EUV channels are sensitive to emission from plasma in the 0.1 to 0.7 MK temperature range, we show that all of AIA’s EUV channels will brighten simultaneously during flares like this, in which energy transport is dominated by nonthermal particle beams. The 2010 July 31 flare observed by AIA and *RHESSI* displays this behavior, and is consistent with the idea that particle beams transported the energy from the energy release site to the chromosphere, where it began to drive chromospheric evaporation nearly 6 minutes before flare temperatures around 10 MK were reached. When thermal conduction from a directly-heated, hot (~ 10 MK) plasma is the dominant energy transport mechanism, the AIA channels that are sensitive to emission from such temperatures (particularly the 94 and 131 Å channels) will brighten earlier than the channels that are not sensitive to such temperatures (171 and 211 Å). Thus, based on the differences between AIA’s response to flares whose energy transport is dominated by nonthermal particle beams from those whose energy transport is dominated by thermal conduction, AIA can be used to determine the dominant energy transport mechanism for any given event.

*Acknowledgements.* J. W. B. acknowledges NASA support through SR&T grant NNX10AC08G. G. D. H. acknowledges partial support from SR&T grant NNX10AC08G and the *RHESSI* Project. The AIA data used are provided courtesy of NASA/SDO and the AIA science team.

## References

Antonucci, E., Alexander, D., Culhane, J. L., de Jager, C., MacNeice, P., Somov, B. V., & Zarro, D. M. 1999, in *The Many Faces of the Sun*, ed. K. T. Strong, J. L. R. Saba, B. M. Haisch, & J. T. Schmelz (New York: Springer), p. 331

Antonucci, E., & Dennis, B. R. 1983, *Sol. Phys.*, 86, 67

Antonucci, E., et al. 1982, *Sol. Phys.*, 78, 107

Aschwanden, M. J., & Boerner, P. 2011, *ApJ*, 732, 81

Boerner, P., et al. 2011, *Sol. Phys.*, DOI: 10.1007/s11207-011-9804-8

Bormmann, P. L. 1999, in *The Many Faces of the Sun*, ed. K. T. Strong, J. L. R. Saba, B. M. Haisch, & J. T. Schmelz (New York: Springer), p. 301

Brosius, J. W. 2003, *ApJ*, 586, 1417

Brosius, J. W. 2009, *ApJ*, 701, 1209

Brosius, J. W., & Holman, G. D. 2007, *ApJ*, 659, L73

Brosius, J. W., & Holman, G. D. 2009, *ApJ*, 692, 492

Brosius, J. W., & Holman, G. D. 2010, *ApJ*, 720, 1472

Brosius, J. W., & Phillips, K. J. H. 2004, *ApJ*, 613, 580

Brosius, J. W., Rabin, D. M., Thomas, R. J., & Landi, E. 2008, *ApJ*, 677, 781

Canfield, R. C., Zarro, D. M., Metcalf, T. R., & Lemen, J. R. 1990, *ApJ*, 348, 333

Culhane, J. L., et al. 2007, *Sol. Phys.*, 243, 19

Delaboudinière, J.-P., et al. 1995, *Sol. Phys.*, 162, 291

Del Zanna, G., O’Dwyer, B., & Mason, H. E. 2011, *A&A*, 535, A46

Dere, K. P., & Cook, J. W. 1979, *ApJ*, 229, 772

Dere, K. P., Landi, E., Mason, H. E., Monsignor Fossi, B. C., & Young, P. R. 1997, *A&AS*, 125, 149

Dere, K. P., Landi, E., Young, P. R., Del Zanna, G., Landini, M., & Mason, H. E. 2009, *A&A*, 498, 915

Doschek, G. A. 1990, *ApJS*, 73, 117

Falewicz, R., Siarkowski, M., & Rudawy, P. 2011, *ApJ*, 733, 37

Fisher, G. H., Canfield, R. C., & McClymont, A. N. 1985a, *ApJ*, 289, 414

Fisher, G. H., Canfield, R. C., & McClymont, A. N. 1985b, *ApJ*, 289, 425

Fisher, G. H., Canfield, R. C., & McClymont, A. N. 1985c, *ApJ*, 289, 434

Fludra, A., Lemen, J. R., Jakimiec, J., Bentley, R. D., & Sylwester, J. 1989, *ApJ*, 344, 991

Harrison, R. A., et al. 1995, *Sol. Phys.*, 162, 233

Hurford, G. J., Schmah, E. J., Schwartz, R. A., et al. 2002, *Sol. Phys.*, 210, 61

Kelly, R. L. 1987, *J. Phys. Chem. Ref. Data*, 16 (Finding List)

Lemen, J. R., et al. 2011, *Sol. Phys.*, DOI: 10.1007/s11207-011-9776-8

Lin, R. P., et al. 2002, *Sol. Phys.*, 210, 3

Mariska, J. T. 1994, *ApJ*, 434, 756

Mariska, J. T., Doschek, G. A., & Bentley, R. D. 1993, *ApJ*, 419, 418

Milligan, R. O., & Dennis, B. R. 2009, *ApJ*, 699, 968

Neupert, W. M. 1968, *ApJ*, 153, L59

O’Dwyer, B., Del Zanna, G., Mason, H. E., Weber, M. A., & Tripathi, D. 2010, *A&A*, 521, A21

Siarkowski, M., Falewicz, R., & Rudawy, P. 2009, *ApJ*, 705, L143

Silva, A. V. R., Wang, H., Gary, D. E., Nitta, N., & Zirin, H. 1997, *ApJ*, 481, 978

Smith, D. M., et al. 2002, *Sol. Phys.*, 210, 33

Wang, T. J., Brosius, J. W., Thomas, R. J., Rabin, D. M., & Davila, J. M. 2010, *ApJS*, 186, 222

Wang, T. J., Thomas, R. J., Brosius, J. W., Young, P. R., Rabin, D. M., Davila, J. M., & Del Zanna, G. 2011, *ApJS*, 197, in press

Zarro, D. M., Canfield, R. C., Strong, K. T., & Metcalf, T. R. 1988, *ApJ*, 324, 582